

How Sensitive is the CMB to a Single Lens?

Ben Rathaus

Tel-Aviv University

September 29, 2011

arXiv:1105.2940: B. Rathaus, A. Fialkov and N. Itzhaki
JCAP **1106**, 033 (2011)

Motivation

- Λ CDM is statistically isotropic.
- Within this framework, weak lensing (WL) is well understood.
- Nevertheless, there exist many anomalies

Motivation

- Λ CDM is statistically isotropic.
- Within this framework, weak lensing (WL) is well understood.
- Nevertheless, there exist many anomalies

Anomaly	Reference
“Axis of evil”	Tegmark et al , Schwarz et al
Anomalously large bulk flow	Watkins et al, Feldman et al
WMAP cold spot	Cruz et al
Giant rings	Kovetz et al

Motivation

- Λ CDM is statistically isotropic.
- Within this framework, weak lensing (WL) is well understood.
- Nevertheless, there exist many anomalies

Anomaly	Reference
“Axis of evil”	Tegmark et al , Schwarz et al
Anomalously large bulk flow	Watkins et al, Feldman et al
WMAP cold spot	Cruz et al
Giant rings	Kovetz et al

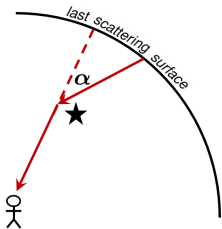
- Can they all be due to an anomalously large structure? Cruz et al, Turok and Spergel, Kovetz et al, Inoue and Silk, Fialkov et al
- Can we detect such a structure by means of WL?

Introduction

- Our Universe is very nearly Gaussian.
- The CMB undergoes WL (where $\alpha = \nabla\psi$):

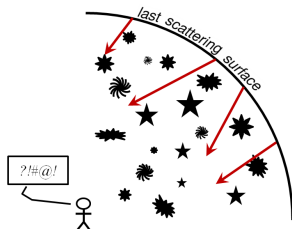
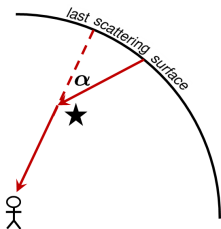
Introduction

- Our Universe is very nearly Gaussian.
- The CMB undergoes WL (where $\alpha = \nabla\psi$):



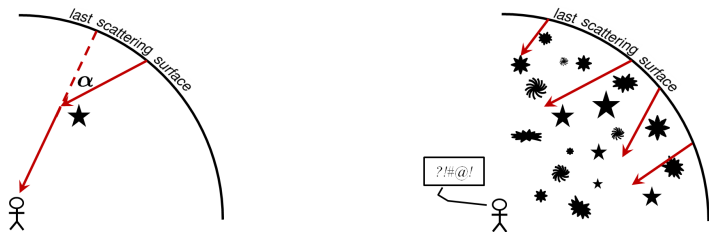
Introduction

- Our Universe is very nearly Gaussian.
- The CMB undergoes WL (where $\alpha = \nabla\psi$):



Introduction

- Our Universe is very nearly Gaussian.
- The CMB undergoes WL (where $\alpha = \nabla\psi$):

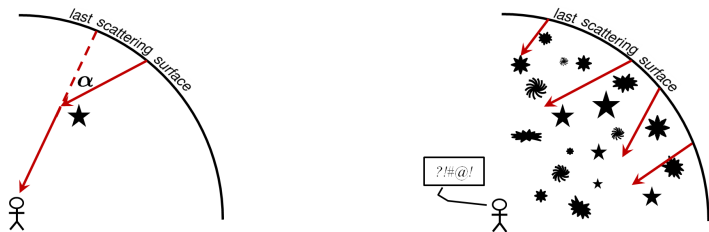


- All we know about the deflection potential

$$\langle \psi(\mathbf{l}) \psi^*(\mathbf{l}') \rangle = (2\pi)^2 \delta(\mathbf{l} - \mathbf{l}') C^\psi(l)$$

Introduction

- Our Universe is very nearly Gaussian.
- The CMB undergoes WL (where $\alpha = \nabla\psi$):



- All we know about the deflection potential

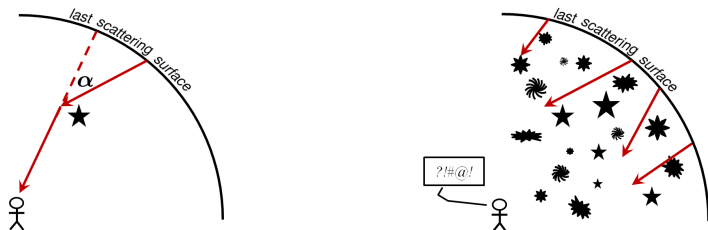
$$\langle \psi(\mathbf{l})\psi^*(\mathbf{l}') \rangle = (2\pi)^2 \delta(\mathbf{l} - \mathbf{l}') C^\psi(l)$$

- The lensed temperature field is also statistically isotropic

$$\langle \tilde{T}(\mathbf{l})\tilde{T}^*(\mathbf{l}') \rangle = (2\pi)^2 \delta(\mathbf{l} - \mathbf{l}') \tilde{C}(l)$$

Introduction

- Our Universe is very nearly Gaussian.
- The CMB undergoes WL (where $\alpha = \nabla\psi$):



- All we know about the deflection potential

$$\langle \psi(\mathbf{l})\psi^*(\mathbf{l}') \rangle = (2\pi)^2 \delta(\mathbf{l} - \mathbf{l}') C^\psi(l)$$

- The lensed temperature field is also statistically isotropic

$$\langle \tilde{T}(\mathbf{l})\tilde{T}^*(\mathbf{l}') \rangle = (2\pi)^2 \delta(\mathbf{l} - \mathbf{l}') \tilde{C}(l)$$

but not entirely Gaussian...

Considering an anomalously large structure

- Such a structure does not respect statistical isotropy.
- Giant structure \longrightarrow deflection potential \longrightarrow deflected T -field.

Considering an anomalously large structure

- Such a structure does not respect statistical isotropy.
- Giant structure \longrightarrow deflection potential \longrightarrow deflected T -field.
- Two *relevant* deformations:
 - Of the deflection potential: $\psi^\Lambda \rightarrow \psi^\Lambda + \delta\psi$.
 - Of the temperature field: $T(\hat{\mathbf{n}}) \rightarrow \tilde{T}(\hat{\mathbf{n}}) = T(\hat{\mathbf{n}} + \nabla\psi)$.
- Both stem from the same structure.
- Both deform one Gaussian field into another Gaussian field.

Considering an anomalously large structure

- Such a structure does not respect statistical isotropy.
- Giant structure \longrightarrow deflection potential \longrightarrow deflected T -field.
- Two *relevant* deformations:
 - Of the deflection potential: $\psi^\Lambda \rightarrow \psi^\Lambda + \delta\psi$.
 - Of the temperature field: $T(\hat{\mathbf{n}}) \rightarrow \tilde{T}(\hat{\mathbf{n}}) = T(\hat{\mathbf{n}} + \nabla\psi)$.
- Both stem from the same structure.
- Both deform one Gaussian field into another Gaussian field.
- *Nevertheless they are essentially different.*

Considering an anomalously large structure

- Such a structure does not respect statistical isotropy.
- Giant structure \longrightarrow deflection potential \longrightarrow deflected T -field.
- Two *relevant* deformations:
 - Of the deflection potential: $\psi^\Lambda \rightarrow \psi^\Lambda + \delta\psi$.
 - Of the temperature field: $T(\hat{\mathbf{n}}) \rightarrow \tilde{T}(\hat{\mathbf{n}}) = T(\hat{\mathbf{n}} + \nabla\psi)$.
- Both stem from the same structure.
- Both deform one Gaussian field into another Gaussian field.
- *Nevertheless they are essentially different.*
- Now construct a S/N ratio.

The deformed deflection potential

Beware!! Hand-waving ahead.

The deformed deflection potential

Beware!! Hand-waving ahead.

- The deformation of the deflection potential is

$$\psi^\Lambda \rightarrow \psi^\Lambda + \delta\psi.$$

The deformed deflection potential

Beware!! Hand-waving ahead.

- The deformation of the deflection potential is

$$\psi^\Lambda \rightarrow \psi^\Lambda + \delta\psi.$$

- The “signal” is the information held in the deformation, $\delta\psi$.
- The “noise” is the background ψ^Λ .
- All we know about ψ^Λ is through its spectrum C^ψ , so

The deformed deflection potential

Beware!! Hand-waving ahead.

- The deformation of the deflection potential is

$$\psi^\Lambda \rightarrow \psi^\Lambda + \delta\psi.$$

- The “signal” is the information held in the deformation, $\delta\psi$.
- The “noise” is the background ψ^Λ .
- All we know about ψ^Λ is through its spectrum C^ψ , so

$$\left(\frac{\mathbf{S}}{\mathbf{N}}\right)^2 = \int \frac{d\mathbf{l}}{(2\pi)^2} \frac{|\delta\psi(\mathbf{l})|^2}{C^\psi(\mathbf{l})}$$

The deformed deflection potential

Beware!! Hand-waving ahead.

- The deformation of the deflection potential is

$$\psi^\Lambda \rightarrow \psi^\Lambda + \delta\psi.$$

- The “signal” is the information held in the deformation, $\delta\psi$.
- The “noise” is the background ψ^Λ .
- All we know about ψ^Λ is through its spectrum C^ψ , so

$$\left(\frac{\mathbf{S}}{\mathbf{N}}\right)^2 = \int \frac{d\mathbf{l}}{(2\pi)^2} \frac{|\delta\psi(l)|^2}{C^\psi(l)}$$

“Ideal S/N”

The ideal S/N

- How come hand-waving was good enough?

$$\langle \psi^\Lambda \rangle = 0 \rightarrow \langle \psi^\Lambda + \delta\psi \rangle = \delta\psi.$$

The ideal S/N

- How come hand-waving was good enough?

$$\langle \psi^\Lambda \rangle = 0 \rightarrow \langle \psi^\Lambda + \delta\psi \rangle = \delta\psi.$$

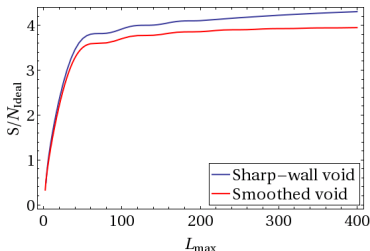
- Though straightforward, difficult to implement.
- Convenient for a given model, as an upper bound.

The ideal S/N

- How come hand-waving was good enough?

$$\langle \psi^\Lambda \rangle = 0 \rightarrow \langle \psi^\Lambda + \delta\psi \rangle = \delta\psi.$$

- Though straightforward, difficult to implement.
- Convenient for a given model, as an upper bound.
- An example: a void.



The realistic S/N

- The observable field of the CMB is the (lensed) temperature field.
- The deformed field gives rise to $S/N_{\text{Realistic}}$.

The realistic S/N

- The observable field of the CMB is the (lensed) temperature field.
- The deformed field gives rise to $S/N_{\text{Realistic}}$.
- A naïve and *wrong* approach would be to follow the ideal case.

The realistic S/N

- The observable field of the CMB is the (lensed) temperature field.
- The deformed field gives rise to $S/N_{\text{Realistic}}$.
- A naïve and *wrong* approach would be to follow the ideal case.
- Deformed field $T \rightarrow \tilde{T}$ belongs to a different type of deformation.

Case	Before deformation	After deformation
Ideal	$\langle \psi \rangle = 0; \quad \langle \psi \psi^* \rangle = C^\psi$	$\langle \psi \rangle = \delta\psi; \quad \langle \psi \psi^* \rangle = C^\psi$
Realistic	$\langle T \rangle = 0; \quad \langle TT^* \rangle = C^{(0)}$	$\langle \tilde{T} \rangle = 0; \quad \langle \tilde{T} \tilde{T}^* \rangle = \tilde{C}$

The realistic S/N

- The observable field of the CMB is the (lensed) temperature field.
- The deformed field gives rise to $S/N_{\text{Realistic}}$.
- A naïve and *wrong* approach would be to follow the ideal case.
- Deformed field $T \rightarrow \tilde{T}$ belongs to a different type of deformation.

Case	Before deformation	After deformation
Ideal	$\langle \psi \rangle = 0; \quad \langle \psi \psi^* \rangle = C^\psi$	$\langle \psi \rangle = \delta\psi; \quad \langle \psi \psi^* \rangle = C^\psi$
Realistic	$\langle T \rangle = 0; \quad \langle TT^* \rangle = C^{(0)}$	$\langle \tilde{T} \rangle = 0; \quad \langle \tilde{T} \tilde{T}^* \rangle = \tilde{C}$

- The deformed covariance matrix is

$$\tilde{C} = C^{(0)} + \epsilon C^{(1)} + \frac{\epsilon^2}{2} C^{(2)} + \dots$$

The realistic S/N

- The observable field of the CMB is the (lensed) temperature field.
- The deformed field gives rise to $S/N_{\text{Realistic}}$.
- A naïve and *wrong* approach would be to follow the ideal case.
- Deformed field $T \rightarrow \tilde{T}$ belongs to a different type of deformation.

Case	Before deformation	After deformation
Ideal	$\langle \psi \rangle = 0; \quad \langle \psi \psi^* \rangle = C^\psi$	$\langle \psi \rangle = \delta\psi; \quad \langle \psi \psi^* \rangle = C^\psi$
Realistic	$\langle T \rangle = 0; \quad \langle TT^* \rangle = C^{(0)}$	$\langle \tilde{T} \rangle = 0; \quad \langle \tilde{T} \tilde{T}^* \rangle = \tilde{C}$

- The deformed covariance matrix is

$$\tilde{C} = C^{(0)} + \epsilon C^{(1)} + \frac{\epsilon^2}{2} C^{(2)} + \dots$$

- The S/N is then

$$\left(\frac{S}{N} \right)^2 = \frac{\epsilon^2}{2} \sum \frac{|C_{ij}^{(1)}|^2}{C_{ii}^{(0)} C_{jj}^{(0)}} = \int \frac{d\mathbf{l}}{(2\pi)^2} \frac{d\mathbf{l}'}{(2\pi)^2} \frac{|C^{(1)}(\mathbf{l}, \mathbf{l}')|^2}{C(l)C(l')}.$$

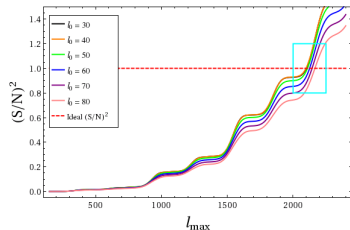
The puzzle

- At most: realistic \rightarrow ideal.

The puzzle

- At most: realistic \rightarrow ideal.
- Ouch!!!
- Gaussian background??

Ideal and Realistic S/N

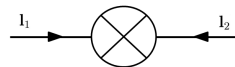


The puzzle

- At most: realistic \rightarrow ideal.
- Ouch!!!
- Gaussian background??
- Our approach: Feynman rules.

Feynman rules

- $T =$ propagating field
- Each leg:
 - 2D momentum
 - propagator $C(l)$
- Integrate loop mom.
- Single lens vertex



$$= \tilde{\gamma}(\mathbf{l}_1, \mathbf{l}_2)/2$$

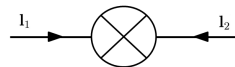
The puzzle

- At most: realistic \rightarrow ideal.
- Ouch!!!
- Gaussian background??
- Our approach: Feynman rules.
- Then the S/N

$$\left(\frac{S}{N}\right)^2 = \int \frac{d\mathbf{l}}{(2\pi)^2} \frac{d\mathbf{l}'}{(2\pi)^2} \frac{|C^{(1)}(\mathbf{l}, \mathbf{l}')|^2}{C(\mathbf{l})C(\mathbf{l}'')}$$

Feynman rules

- $T =$ propagating field
- Each leg:
 - 2D momentum
 - propagator $C(l)$
- Integrate loop mom.
- Single lens vertex

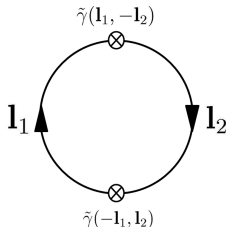


$$= \tilde{\gamma}(\mathbf{l}_1, \mathbf{l}_2)/2$$

The puzzle

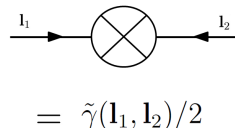
- At most: realistic \rightarrow ideal.
- Ouch!!!
- Gaussian background??
- Our approach: Feynman rules.
- Then the S/N

$$\left(\frac{S}{N}\right)^2 = \int \frac{d\mathbf{l}}{(2\pi)^2} \frac{d\mathbf{l}'}{(2\pi)^2} \frac{|C^{(1)}(\mathbf{l}, \mathbf{l}')|^2}{C(\mathbf{l})C(\mathbf{l}'')}$$



Feynman rules

- $T =$ propagating field
- Each leg:
 - 2D momentum
 - propagator $C(l)$
- Integrate loop mom.
- Single lens vertex



The HEP-way

- Our propagating field is Λ CDM temperature field.
- Gaussianity means

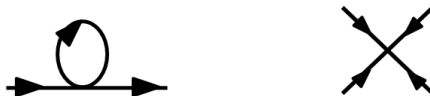


The HEP-way

- Our propagating field is Λ CDM temperature field.
- Gaussianity means



- Non-Gaussianity, on the other hand may mean



The HEP-way

- Our propagating field is Λ CDM temperature field.
- Gaussianity means



- Non-Gaussianity, on the other hand may mean



and this is indeed the case in Λ CDM WL.

How does this come about?

- First, write a 4-point function in terms of the primordial temperature field.
- Second, keep only second (lowest) order in the deflection potential

$$\begin{aligned}
 \langle \tilde{T} \tilde{T} \tilde{T} \tilde{T} \rangle &\rightarrow \langle (T + \nabla \psi \cdot \nabla T) (T + \nabla \psi \cdot \nabla T) T T \rangle \\
 &\rightarrow \langle (\nabla \psi \cdot \nabla T) (\nabla \psi \cdot \nabla T) T T \rangle
 \end{aligned}$$

How does this come about?

- First, write a 4-point function in terms of the primordial temperature field.
- Second, keep only second (lowest) order in the deflection potential

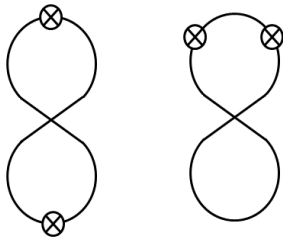
$$\begin{aligned} \langle \tilde{T} \tilde{T} \tilde{T} \tilde{T} \rangle &\rightarrow \langle (T + \nabla \psi \cdot \nabla T) (T + \nabla \psi \cdot \nabla T) T T \rangle \\ &\rightarrow \langle (\nabla \psi \cdot \nabla T) (\nabla \psi \cdot \nabla T) T T \rangle \end{aligned}$$

- Finally, consider, for example, the following contraction

$$\langle (\nabla \psi \cdot \nabla T) (\nabla \psi \cdot \nabla T) T T \rangle$$

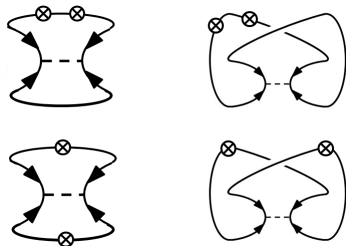
Resolving the puzzle

- 2-loop corrections to the S/N.



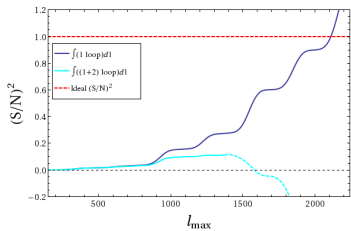
Resolving the puzzle

- 2-loop corrections to the S/N.
- Moreover, 4-point vertex has substructure.



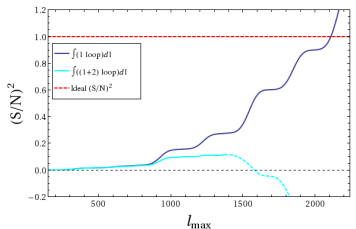
Resolving the puzzle

- 2-loop corrections to the S/N.
- Moreover, 4-point vertex has substructure.
- These corrections are *negative* therefore reduce the S/N.



Resolving the puzzle

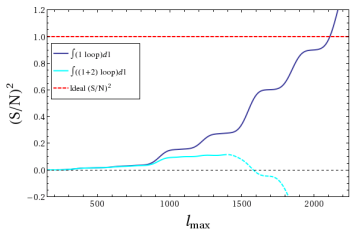
- 2-loop corrections to the S/N.
- Moreover, 4-point vertex has substructure.
- These corrections are *negative* therefore reduce the S/N.
- S/N turn-over is where 3-loop come into play (the latest).



Resolving the puzzle

- 2-loop corrections to the S/N.
- Moreover, 4-point vertex has substructure.
- These corrections are *negative* therefore reduce the S/N.
- S/N turn-over is where 3-loop come into play (the latest).
- We find the relation

$$\left(\frac{S}{N}\right)_{\text{Accessible}} \sim \frac{1}{\sqrt{10}} \left(\frac{S}{N}\right)_{\text{Ideal}}$$

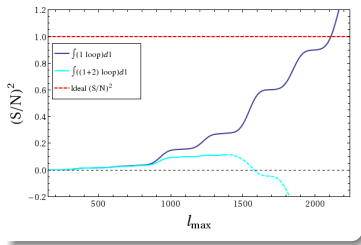


Resolving the puzzle

- 2-loop corrections to the S/N.
- Moreover, 4-point vertex has substructure.
- These corrections are *negative* therefore reduce the S/N.
- S/N turn-over is where 3-loop come into play (the latest).
- We find the relation

$$\left(\frac{S}{N}\right)_{\text{Accessible}} \sim \frac{1}{\sqrt{10}} \left(\frac{S}{N}\right)_{\text{Ideal}}$$

- WMAP Vs. Planck



Resolving the puzzle

- 2-loop corrections to the S/N.
- Moreover, 4-point vertex has substructure.
- These corrections are *negative* therefore reduce the S/N.
- S/N turn-over is where 3-loop come into play (the latest).
- We find the relation

$$\left(\frac{S}{N}\right)_{\text{Accessible}} \sim \frac{1}{\sqrt{10}} \left(\frac{S}{N}\right)_{\text{Ideal}}$$

- WMAP Vs. Planck

Example

Cosmic void

- WMAP Cold-spot
- Argued: S/N ~ 40
- *But:*
 - S/N_{Ideal} ~ 4
 - S/N_{Acc.} ~ 1

Thank You!